

Direct detection of relic active and sterile neutrinos

Yu-Feng Li

Institute of High Energy Physics, Chinese Academy of Sciences, P.O. Box 918, Beijing 100049, China

E-mail: liyufeng@ihep.ac.cn

Abstract. Both active and sterile sub-eV neutrinos can form the cosmic neutrino background in the early Universe. We consider the beta-decaying (e.g., ${}^3\text{H}$) and EC-decaying (e.g., ${}^{163}\text{Ho}$) nuclei as the promising targets to capture relic neutrinos in the laboratory. We calculate the capture rates of relic electron neutrinos and antineutrinos against the corresponding beta decay or electron capture (EC) decay backgrounds in the $(3+N_s)$ flavor mixing scheme, and discuss the future prospect in terms of the PTOLEMY project. We stress that such direct measurements of hot DM might not be hopeless in the long term.

1. Introduction

Both active and sterile sub-eV neutrinos can form the cosmic neutrino background ($\text{C}\nu\text{B}$) when they were decoupled from radiation and matter at a temperature of about one MeV and an age of one second after the Big Bang [1]. Relic neutrinos played important roles in the evolution of the Universe, and has been indirectly proved from cosmological data on the Big Bang nucleosynthesis (BBN), cosmic microwave background (CMB) anisotropies and large-scale structures (LSS) [2]. Without considering the lepton asymmetries, the temperature and average number density for one species of relic neutrinos can be expressed as

$$T_\nu = \left(\frac{4}{11}\right)^{1/3} T_\gamma \approx 1.945 \text{ K}, \quad n_\nu = \frac{3}{11} n_\gamma \approx 112 \text{ cm}^{-3}. \quad (1)$$

As a consequence, one predicts the average three-momentum today for each species of the relic neutrino is very small, i.e., $\langle p_\nu \rangle = 3T_\nu \approx 5 \times 10^{-4} \text{ eV}$ [3].

Cosmological observations provide indirect evidence for the existence of the $\text{C}\nu\text{B}$, however, it is a great challenge to the present experimental techniques for the direct detection in a laboratory experiment. Among several detection possibilities [3], the most promising one seems to be the neutrino capture experiment using radioactive β -decaying nuclei [4, 5, 6, 7, 8, 9, 10, 11, 12, 13]. The PTOLEMY project [14] aims to detect the $\text{C}\nu\text{B}$ using 100 grams of ${}^3\text{H}$ as the capture target. Other interesting methods include the electron-capture (EC) decaying nuclei [15, 16, 17, 18], the annihilation of extremely high-energy cosmic neutrinos at the Z -resonance [19, 20, 21], and the atomic de-excitation method [22].

The remaining parts of this work are organized as follows. In Sec. 2 we introduce methods of relic neutrino captures on the beta-decaying and EC-decaying nuclei, and calculate the rates and energy spectra of neutrino capture rates. Sec. 3 is devoted to the flavor effects of relic neutrino captures including the neutrino mass hierarchy and presence of sterile neutrinos, and then conclude in Sec. 4.

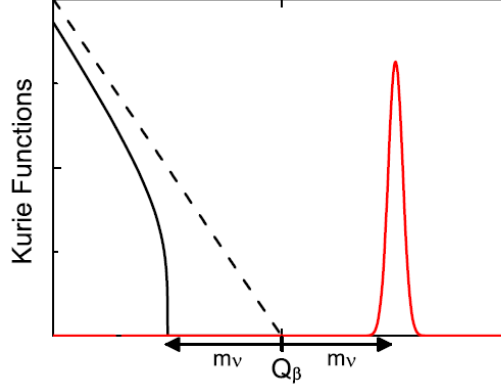


Figure 1. Idealized electron spectra for the tritium beta decay and relic neutrino capture. The dashed and black-solid lines are shown for β -decay spectra of the massless and massive neutrinos respectively. The red-solid line with the sharp peak is for the relic neutrino signal.

2. Captures on Beta-decaying or EC-decaying Nuclei

In the presence of $3+N_s$ species of active and sterile neutrinos, the flavor eigenstates of three active neutrinos and N_s sterile neutrinos can be written as [1, 2]

$$\begin{pmatrix} \nu_e \\ \nu_\mu \\ \nu_\tau \\ \vdots \end{pmatrix} = \begin{pmatrix} U_{e1} & U_{e2} & U_{e3} & \cdots \\ U_{\mu 1} & U_{\mu 2} & U_{\mu 3} & \cdots \\ U_{\tau 1} & U_{\tau 2} & U_{\tau 3} & \cdots \\ \vdots & \vdots & \vdots & \ddots \end{pmatrix} \begin{pmatrix} \nu_1 \\ \nu_2 \\ \nu_3 \\ \vdots \end{pmatrix}, \quad (2)$$

where ν_i is a mass eigenstate of active (for $1 \leq i \leq 3$) or sterile (for $4 \leq i \leq 3 + N_s$) neutrinos, and $U_{\alpha i}$ stands for an element of the $(3 + N_s) \times (3 + N_s)$ neutrino mixing matrix.

For the nuclear β -decay process with the mass number A and atomic number Z of the parent nucleus, i.e. $\mathcal{N}(A, Z) \rightarrow \mathcal{N}'(A, Z + 1) + e^- + \bar{\nu}_e$, the differential decay rate of a β -decay can be written as [23]

$$\begin{aligned} \frac{d\lambda_\beta}{dT_e} &= \int_0^{Q_\beta - \min(m_i)} dT'_e \left\{ \frac{G_F^2 \cos^2 \theta_C}{2\pi^3} F(Z, E_e) |\mathcal{M}|^2 E_e \sqrt{E_e^2 - m_e^2} \right. \\ &\times (Q_\beta - T'_e) \sum_{i=1}^{3+N_s} \left[|U_{ei}|^2 \sqrt{(Q_\beta - T'_e)^2 - m_i^2} \Theta(Q_\beta - T'_e - m_i) \right] \Big\} \times R(T_e, T'_e), \quad (3) \end{aligned}$$

where $T'_e = E_e - m_e$ denotes the kinetic energy of the outgoing electron, $F(Z, E_e)$ is the Fermi function, $|\mathcal{M}|^2$ is the dimensionless nuclear matrix elements [23], and $\theta_C \simeq 13^\circ$ is the Cabibbo angle. In addition, a Gaussian energy resolution function

$$R(T_e, T'_e) = \frac{1}{\sqrt{2\pi} \sigma} \exp \left[-\frac{(T_e - T'_e)^2}{2\sigma^2} \right], \quad (4)$$

is implemented in Eq. (3) to include the finite energy resolution, and the theta function $\Theta(Q_\beta - T'_e - m_i)$ is adopted to ensure the kinematic requirement. The spectral shape near the β -decay endpoint represents a kinetic measurement of the absolute neutrino masses, which can be understood by comparing the dashed and black solid lines of Fig. 1.

On the other hand, the threshold-less neutrino capture process,

$$\nu_e + \mathcal{N}(A, Z) \rightarrow \mathcal{N}'(A, Z + 1) + e^-, \quad (5)$$

is located well beyond the end point of the β -decay, where the signal is characterized by the monoenergetic kinetic energy of the electron for each mass eigenstate (see the red-solid line in Fig. 1). This capture process is suitable to detect relic active and sterile neutrinos, and a measurement of the distance between the decay and capture processes will directly probe the $C\nu B$. The differential neutrino capture rate of this process reads

$$\frac{d\lambda_\nu}{dT_e} = \sum_i |U_{ei}|^2 \sigma_{\nu_i} v_{\nu_i} n_{\nu_i} R(T_e, T_e^i), \quad (6)$$

where the sum is for all the neutrino mass eigenstates and $n_{\nu_i} \equiv \zeta_i \cdot \langle n_{\nu_i} \rangle$ denotes the number density of the relic neutrinos ν_i around the Earth. The standard Big Bang cosmology predicts $\langle n_{\nu_i} \rangle \approx \langle n_{\bar{\nu}_i} \rangle \approx 56 \text{ cm}^{-3}$ for each species of active neutrinos, and it is also expected to hold for each sterile neutrino species if they could be fully thermalized in the early Universe. The number density of relic active and sterile neutrinos may be enhanced by the gravitational clustering effect (i.e., the factor ζ_i) when the neutrino mass is greater than 0.1 eV [24]. The capture cross-section times the neutrino velocity can be written as $\sigma_{\nu_i} v_{\nu_i} = 2\pi^2 \ln 2 / (A \times T_{1/2})$, where A is the nuclear factor characterized by Q_β and Z and $T_{1/2}$ is the half-life of the parent nucleus.

To get a better signal-to-background ratio, one can investigate different kinds of candidate nuclei by considering factors including the cross-section, half-life, β -decay rate, and detector energy resolution. Based on this selection criterion, several promising nuclei such as ^3H , ^{106}Ru , and ^{187}Re are identified after an exhaustive survey in Ref. [6].

The β -decay experiments of current generation include the spectrometer of KATRIN [25] and the calorimeter of MARE [26]. KATRIN uses 50 μg of ^3H as the effective target mass, and MARE is planning to deploy 760 grams of ^{187}Re . Therefore, we can estimate their respective $C\nu B$ event rates to be 10^{-6} yr^{-1} and 10^{-7} yr^{-1} without considering the gravitational clustering effect. A first realistic proposal for the $C\nu B$ detection is the PTOLEMY project [14], which is designed to employ 100 grams of ^3H as the capture target using a combination of a large-area surface-deposition tritium target, the MAC-E filter, the RF tracking, the time-of-flight systems, and the cryogenic calorimetry. Finally, the event rate of PTOLEMY are calculated to reach the observable level:

$$N^\nu(\text{PTOLEMY}) \simeq 8.0 \times \sum_i |U_{ei}|^2 \zeta_i \text{ yr}^{-1}. \quad (7)$$

According to Eq. (5), only electron neutrinos can be captured in the β -decaying nuclei. One should consider other possibilities for the cosmic antineutrino background detection. Similar to the process of captures on β -decaying nuclei, the EC-decaying nuclei can be the target of relic antineutrino captures. The isotope ^{163}Ho is a promising candidate in this respect [15, 16, 17, 18]. The properties of the relic antineutrino capture against the EC-decaying background are similar to those of β -decaying nuclei [17]. As the order of magnitude estimate, one needs 30 kg ^{163}Ho to obtain one event per year for the relic antineutrino detection.

3. Flavor Effects

Besides the total capture rates, the $C\nu B$ detection exhibits interesting properties of flavor effects due to the neutrino mixing. In this section, we shall discuss the effects of the neutrino mass hierarchy [8] and presence of light sterile neutrinos [10].

Fig. 2 shows the capture rate of the $C\nu B$ as a function of the kinetic energy T_e of electrons in the standard three-neutrino scheme with $\Delta m_{31}^2 > 0$ and $\Delta m_{31}^2 < 0$. The gravitational clustering of three active neutrinos has been neglected for simplicity. Δ (i.e., $\Delta = 2\sqrt{2 \ln 2} \sigma$) denotes the finite energy resolution. As the lightest neutrino mass increases from 0 to 0.1 eV, the neutrino capture signal moves towards the larger T_e region. The distance between the signal peak and

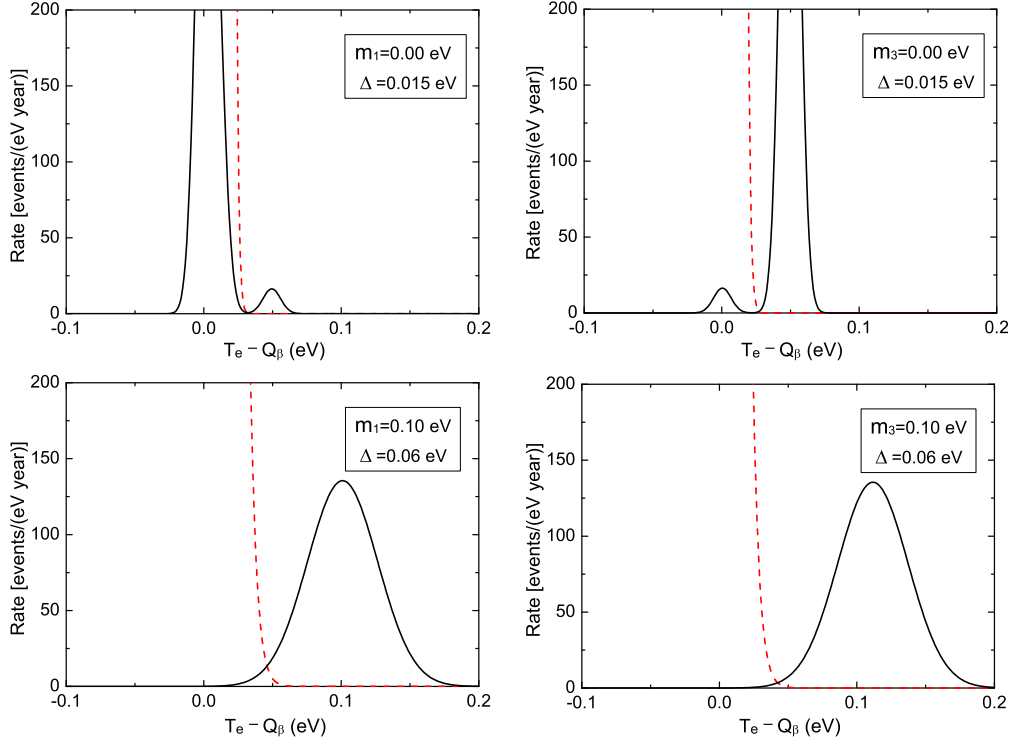


Figure 2. The relic neutrino capture rate as a function of the kinetic energy of electrons in the standard scheme with $\Delta m_{31}^2 > 0$ (left panel) or $\Delta m_{31}^2 < 0$ (right panel). The gravitational clustering of three active neutrinos has been neglected for simplicity. We adopt 100 grams of ${}^3\text{H}$, and best-fit values of the relevant three-neutrino oscillation parameters from Ref. [2].

the β -decay background becomes larger for a larger lightest neutrino mass, and therefore the required energy resolution is less stringent. Comparing between the left panel and right panel, one can observe that it is easier to detect the $C\nu\text{B}$ in the $\Delta m_{31}^2 < 0$ case, where the capture signal is separated more apparently from the β -decay background. The reason is that the dominant mass eigenstates ν_1 and ν_2 in ν_e have greater eigenvalues than in the $\Delta m_{31}^2 > 0$ case.

Next we shall study the (3+2) mixing scheme with two light sterile neutrinos. Considering the hints of short baseline oscillations [27, 28], we assume $m_4 = 0.2$ eV and $m_5 = 0.4$ eV together with $|U_{e1}| \approx 0.792$, $|U_{e2}| \approx 0.534$, $|U_{e3}| \approx 0.168$, $|U_{e4}| \approx 0.171$ and $|U_{e5}| \approx 0.174$ in the numerical calculations. We illustrate the capture rate of the $C\nu\text{B}$ against the corresponding β -decay background for both $\Delta m_{31}^2 > 0$ and $\Delta m_{31}^2 < 0$ schemes in Fig. 3. To take account of possible gravitational clustering effects, we assume $\zeta_1 = \zeta_2 = \zeta_3 = 1$ and $\zeta_5 = 2\zeta_4 = 10$. As one can see from Fig. 3, the signals of sterile neutrinos are obviously enhanced because of $\zeta_4 > 1$ and $\zeta_5 > 1$. If the overdensity of relic neutrinos is very significant around the Earth, it will be helpful for the $C\nu\text{B}$ detection through the neutrino capture process.

4. Conclusion

The standard Big Bang cosmology predicts the existence of a cosmic neutrino background formed at an age of one second after the Big Bang. A direct measurement of the relic neutrinos would open a new window to the early Universe. We have discussed the future prospect for the direct detection of the $C\nu\text{B}$, with the emphasis on the method of captures on β -decaying nuclei and PTOLEMY project. We calculated the neutrino capture rate against the corresponding β -decay background, and discussed the possible flavor effects including the neutrino mass hierarchy and

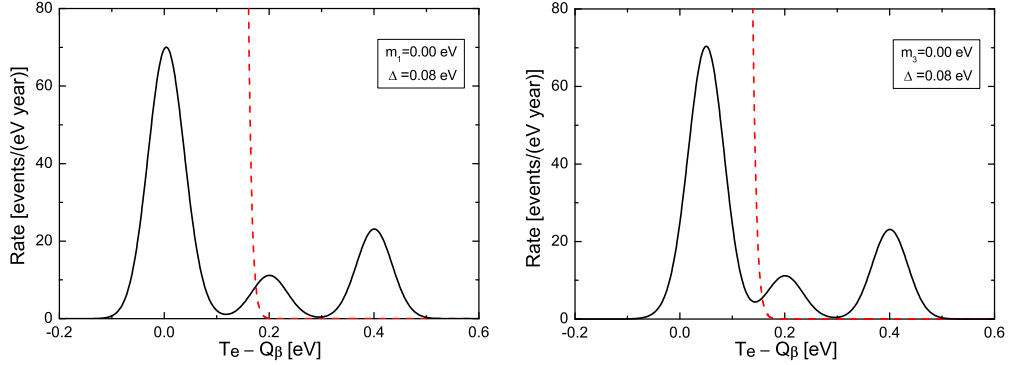


Figure 3. The capture rate of the $C\nu B$ as a function of the electron's kinetic energy in the (3+2) mixing scheme with $\Delta m_{31}^2 > 0$ (left panel) and $\Delta m_{31}^2 < 0$ (right panel) [10]. The gravitational clustering of relic sterile neutrinos around the Earth has been illustrated by taking $\zeta_1 = \zeta_2 = \zeta_3 = 1$ and $\zeta_5 = 2\zeta_4 = 10$.

presence of light sterile neutrinos. We stress that such direct measurements of the $C\nu B$ in the laboratory experiments might not be hopeless in the long term.

Acknowledgement

This work was supported by the National Natural Science Foundation of China under grant Nos. 11135009 and 11305193, and the CAS Center for Excellence in Particle Physics (CCEPP).

References

- [1] See e.g., Xing Z Z and Zhou S 2011 *Neutrinos in Particle Physics, Astronomy and Cosmology* (Berlin: Springer-Verlag)
- [2] Olive K A *et al.* (Particle Data Group) 2014 *Chin. Phys. C* **38** 090001
- [3] For a brief review, see: Ringwald A 2009 *Nucl. Phys. A* **827** 501c
- [4] Weinberg S 1962 *Phys. Rev.* **128** 1457
- [5] Irvine J M and Humphreys R 1983 *J. Phys. G* **9** 847
- [6] Cocco A, Mangano G and Messina M 2007 *JCAP* **0706** 015
- [7] Lazauskas R, Vogel P and Volpe C 2008 *J. Phys. G* **35** 025001
- [8] Blennow M 2008 *Phys. Rev. D* **77** 113014
- [9] Kaboth A, Formaggio J A and Monreal B 2010 *Phys. Rev. D* **82** 062001
- [10] Li Y F, Xing Z Z and Luo S 2010 *Phys. Lett. B* **692** 261
- [11] Li Y F and Xing Z Z 2011 *Phys. Lett. B* **695** 205
- [12] Liao W 2010 *Phys. Rev. D* **82** 073001
- [13] Long A J, Lunardini C and Sabancilar E 2014 *JCAP* **1408** 038
- [14] S. Betts *et al.* (PTOLEMY) 2013 *Preprint* arXiv:1307.4738 [astro-ph]
- [15] Cocco A G, Mangano G and Messina M 2009 *Phys. Rev. D* **79** 053009
- [16] Lusignoli M and Vignati M 2011 *Phys. Lett. B* **697** 11
- [17] Li Y F and Xing Z Z 2011 *Phys. Lett. B* **698** 430
- [18] Li Y F and Xing Z Z 2011 *JCAP* **1108** 006
- [19] Weiler T J 1982 *Phys. Rev. Lett.* **49** 234; 1984 *Astrophys. J.* **285** 495
- [20] Eberle B *et al.* 2004 *Phys. Rev. D* **70** 023007
- [21] Barenboim G, Mena Requejo O and Quigg C 2005 *Phys. Rev. D* **71** 083002
- [22] Yoshimura M, Sasao N and Tanaka M 2015 *Phys. Rev. D* **91** 063516
- [23] Otten E W and Weinheimer C 2008 *Rept. Prog. Phys.* **71** 086201
- [24] Ringwald A and Wong Y Y 2004 *JCAP* **0412** 005
- [25] Osipowicz A *et al.* (KATRIN Collaboration) 2001 *Preprint* arXiv:hep-ex/0109033
- [26] Nucciotti A (MARE Collaboration) 2012 *Nucl. Phys. Proc. Suppl.* **229** 155
- [27] Giunti C *et al.* 2013 *Phys. Rev. D* **88** 073008
- [28] Kopp J *et al.* 2013 *JHEP* **1305** 050